The Cooling of Multi-Stranded Flaring Loops REU Montana State University 2011



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Yohkoh SXT

 I am using data from the Soft Xray Telescope aboard Yohkoh



Solar Flare



Searching for a better model for the cooling of solar flares

time (s)

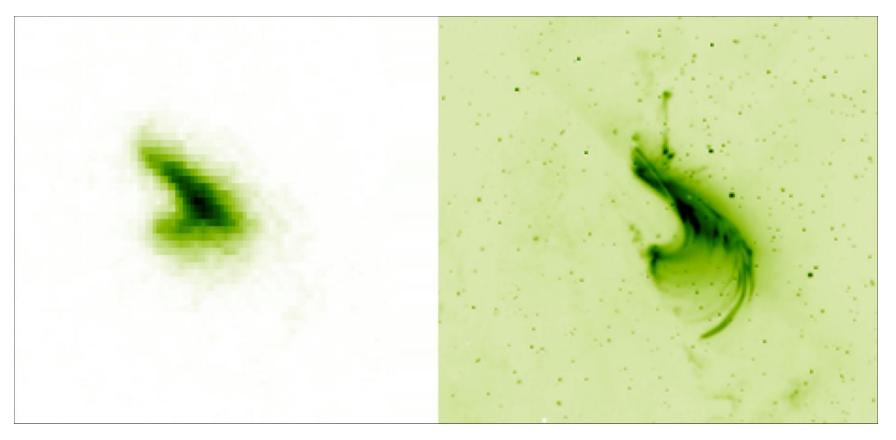
800

time (s)

1000

Current Observing Unresolved Multi-Stranded Loops theories cannot Thin Al Flux $\chi^2 = 8.9e + 03$ Al Mg Flux $\chi^2 = 1.5e + 04$ model the 2.0×10⁶ 1.2×10 1.0×10 cooling very 1.5×10⁶ Flux (DN) 8.0×10 (DN) well. We will 1.0×10⁶ 6.0×10⁵ be working to 4.0×10⁵ 5.0×10⁵ 2.0×10⁵ better fit the tails of these 600 800 1000 1200 time (s) light curves by Thin Be Flux $\chi^2 = 7.4e + 02$ Thick Al Flux $\chi^2 = 9.5e + 02$ 1×10⁵ 6×10 varying 5×10 8×10 parameters 4×10 Tux (DN) DN 6×104 3×10⁴ including lux 4×10 2×10^{4} number of 2×104 1×10⁴ loops. 800 1000 1200 time (s

With better resolution, we can see more loops than were initially visible.

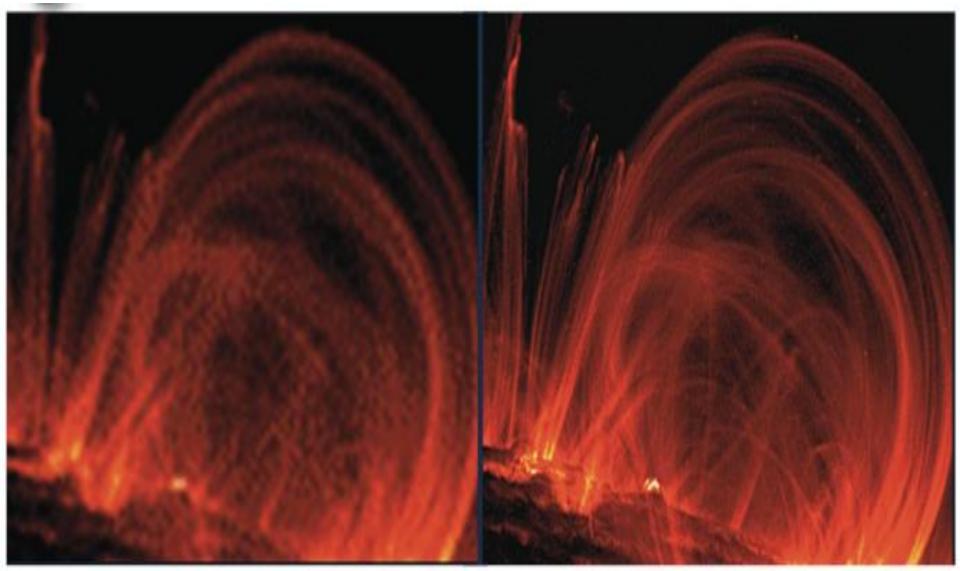


• SXT

TRACE

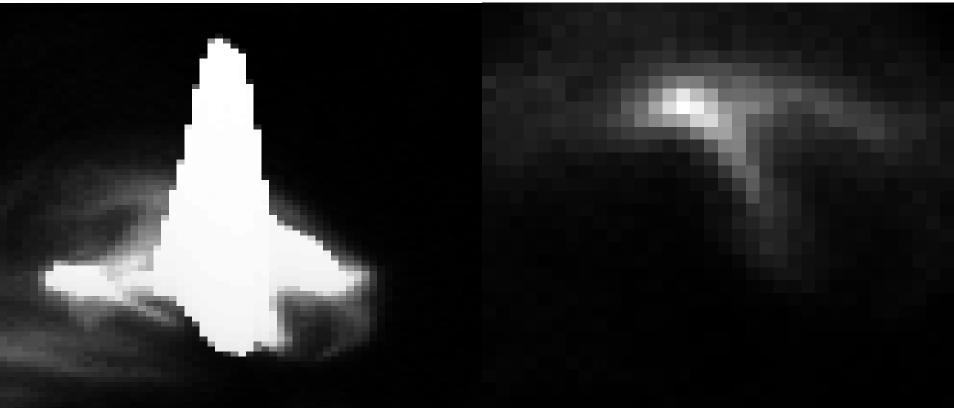
 We believe there are more loops then seen even here that would be visible with even better resolution.

How Could There Be Many Loops and Yet We Don't See Them?



A saturated frame and a normal frame

 When I had retrieved the data from Yohkoh for our project, I had to remove saturated frames like the one below. I also had to remove frames not focused on the flare, scattered light and vignetting.



1992 September 6 Flare Before

• This movie shows what the data looked like before I cleaned it up.



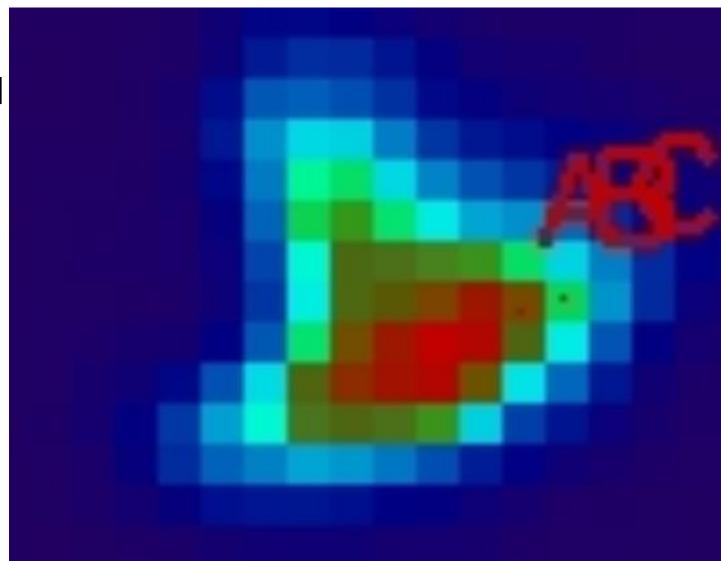
1992, September 6 Flare After

• This movie shows the data after I cleaned it up.



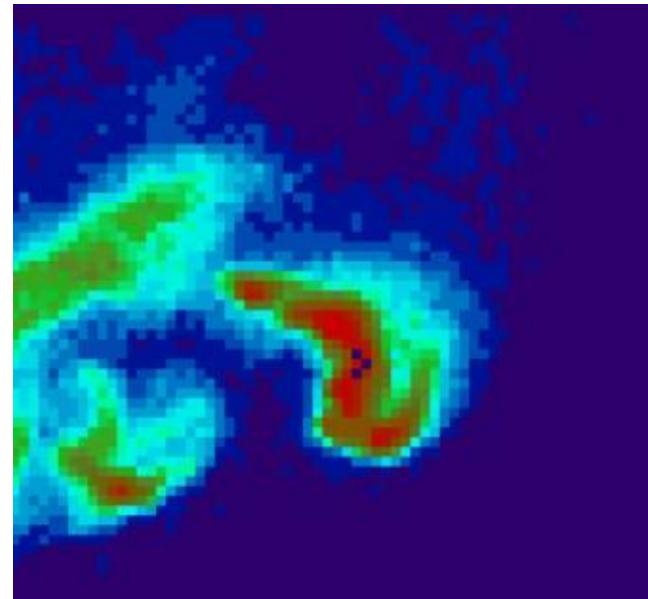
Selected Pixels

You can
 see where I
 picked the
 pixels to
 make light
 curves
 from.

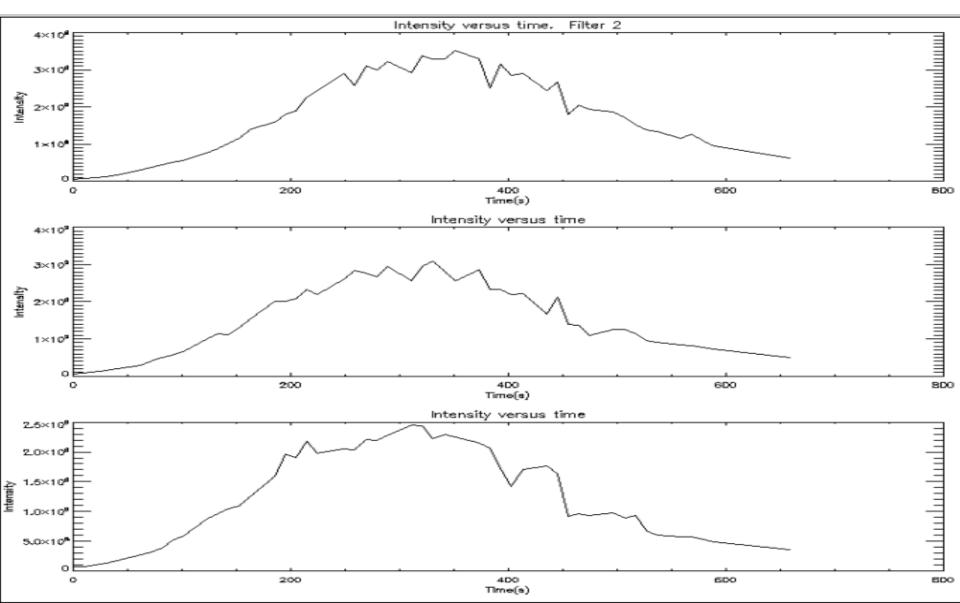


• Here's another view from the movie program we used to help us follow where the pixels were in the flare.

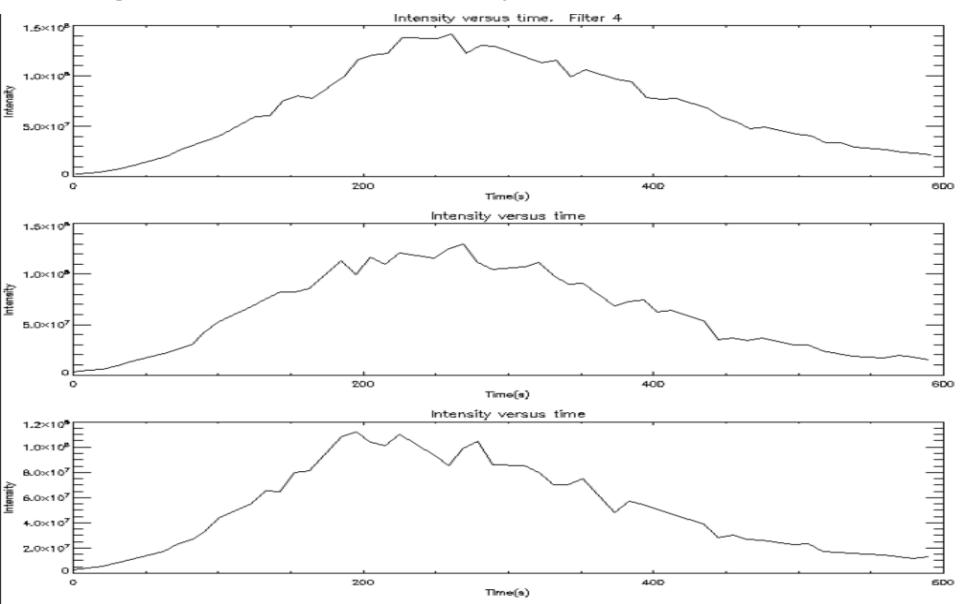
Selected Pixels



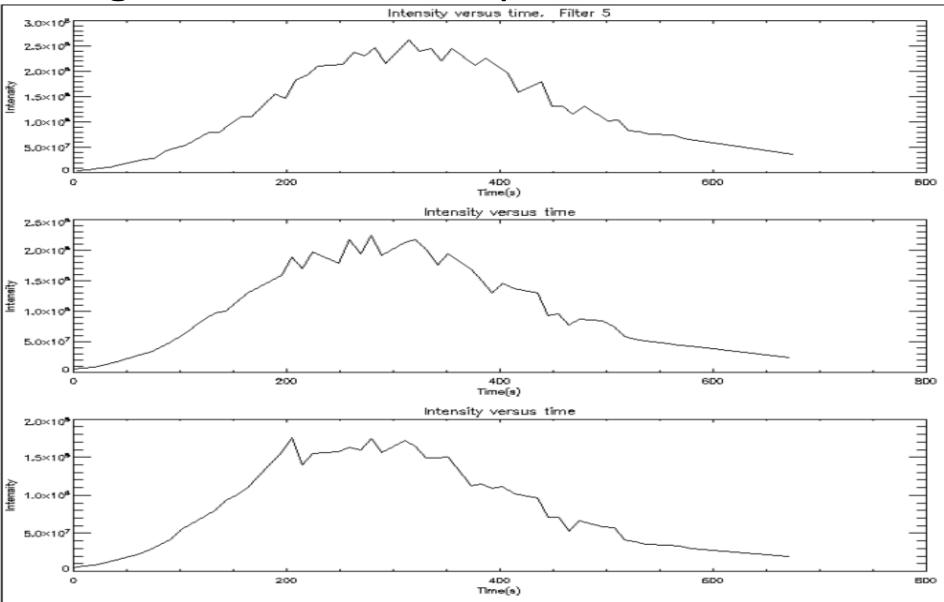
Light Curve 1992 September 6 Filter 2



Light Curve 1992 September 6 Filter 4



Light Curve 1992 September 6 Filter 5



Light Curve Stitching Before Intensity versus time. Filter 5 2.0×10⁸ 1.5×10^e The align ntensity 1.0×10⁸ program 5.0×10⁷ wasn't 0 500 1000 150D 2000 o Time(s) Intensity versus time working 1.5×10⁶ properly. 1.0×10 5.0×10⁷ 0 500 1000 150D 2000 o Time(s) Intensity versus time 20×10 1.5×10⁶ , 1.0×10[®] 1_ 5.0×107

500

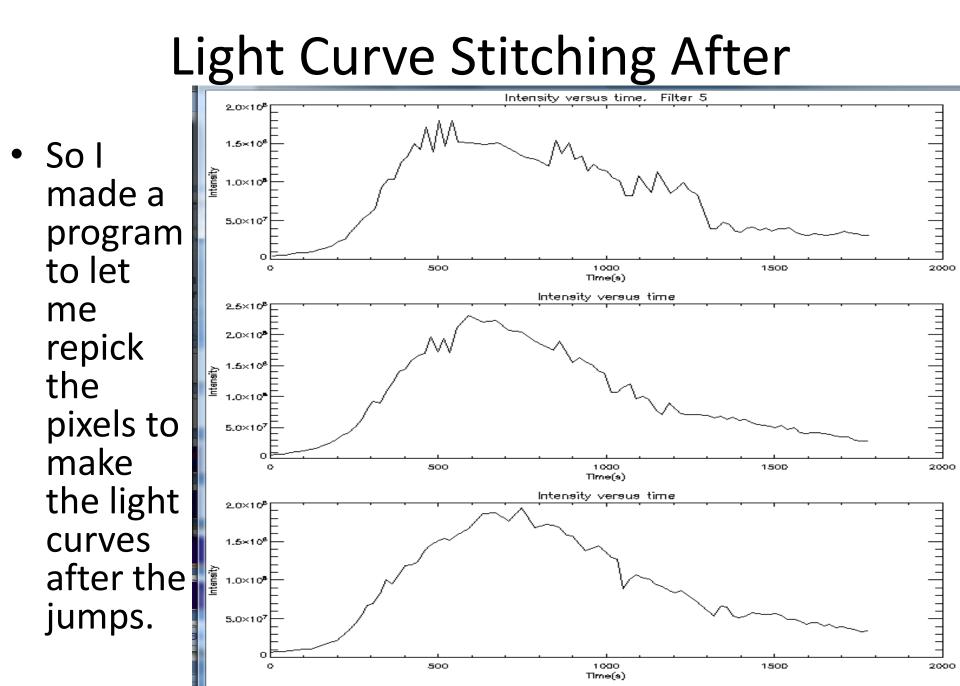
1000

Time(s)

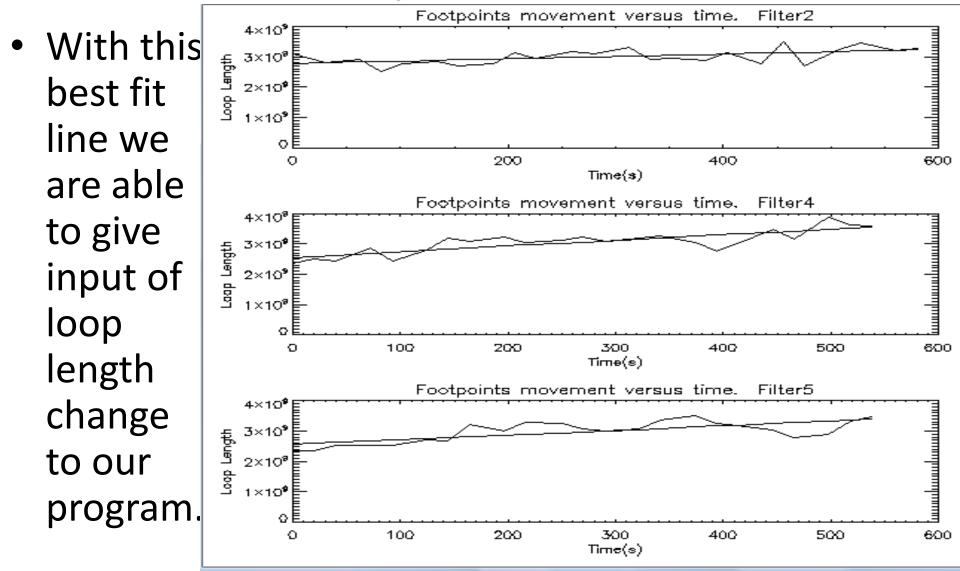
150D

2000

oE



Loop Length vs Time 1992 September 6



Why Does Loop Length Increase?

 Flare strands form on top of older strands. So like this rainbow as you go up the length and foot point distance increases

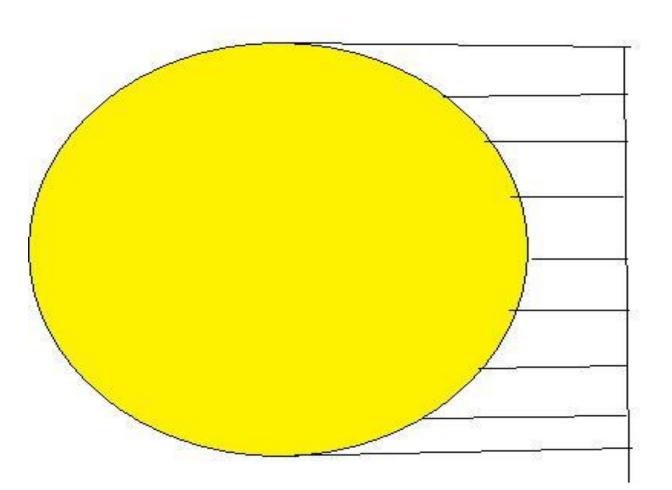
How much do the footpoints move?

- The x's mark the position at 5:14:32.
- The o's mark the position at 5:24:02.



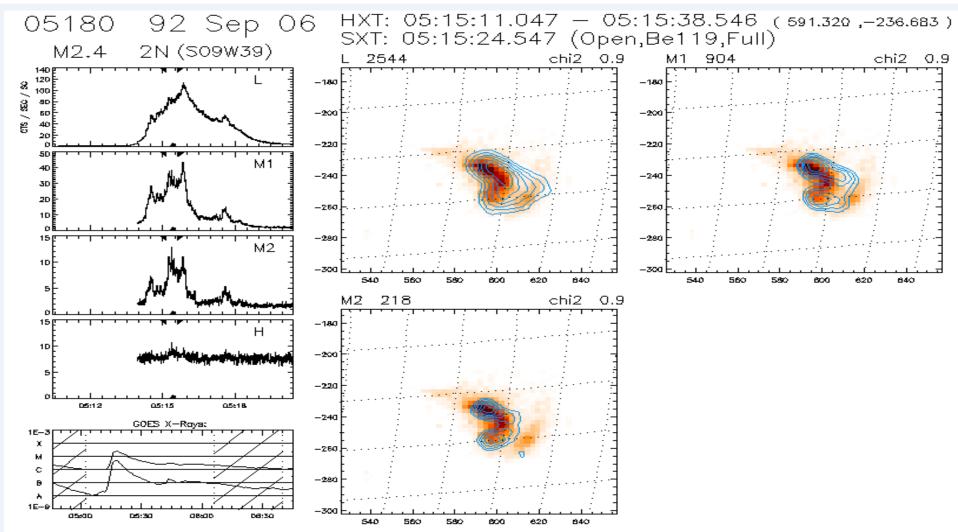
cm per arcsecond on different parts of the sun

 Observing different parts of the sun gives different distances for the same angles, because we are making a 3D object 2D



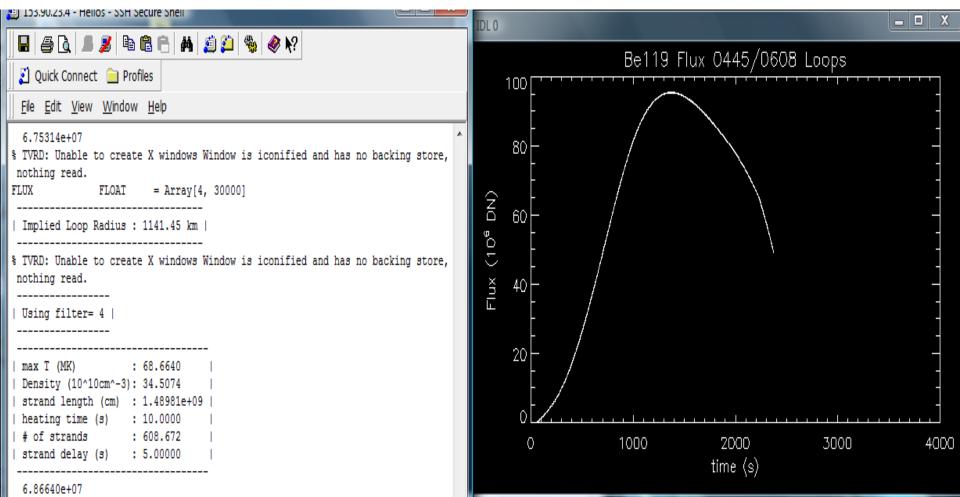
Hard X ray 1992 September 6th

• The contours are the hard x-rays, the colored parts are the soft x-rays

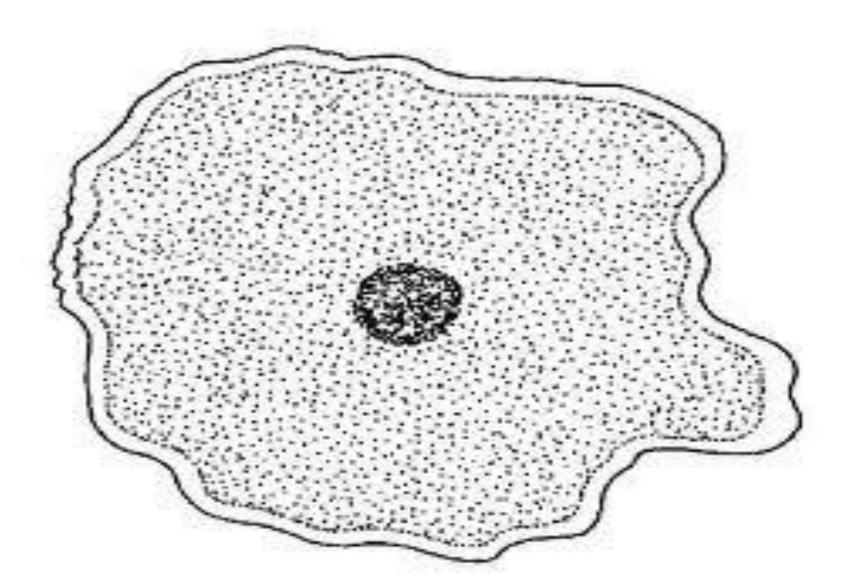


Finding the Best Temperature, Density, and Number of Strands

• This is a picture of the running amoeba program



What is an amoeba?



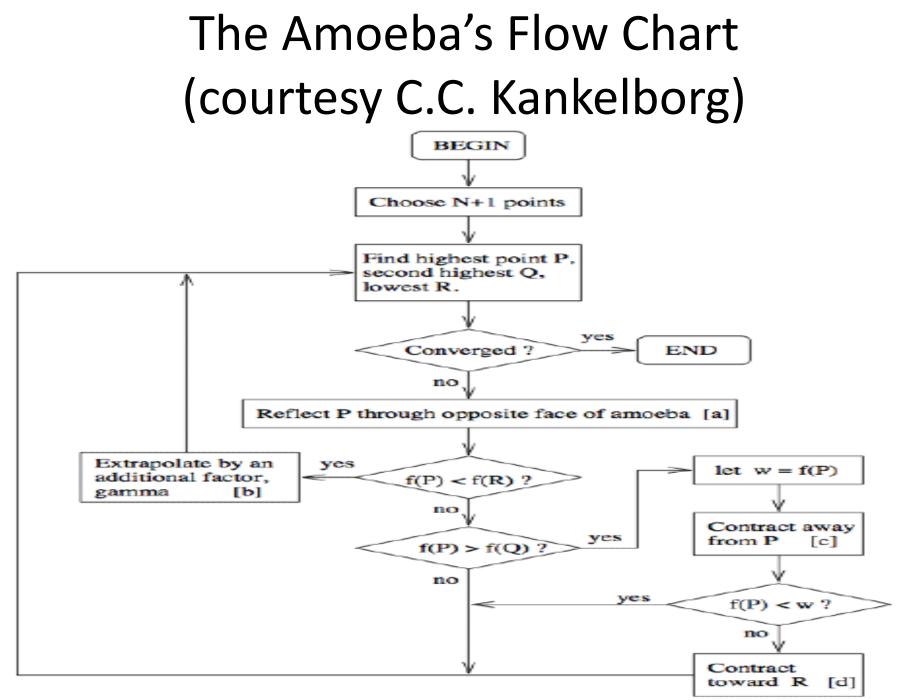
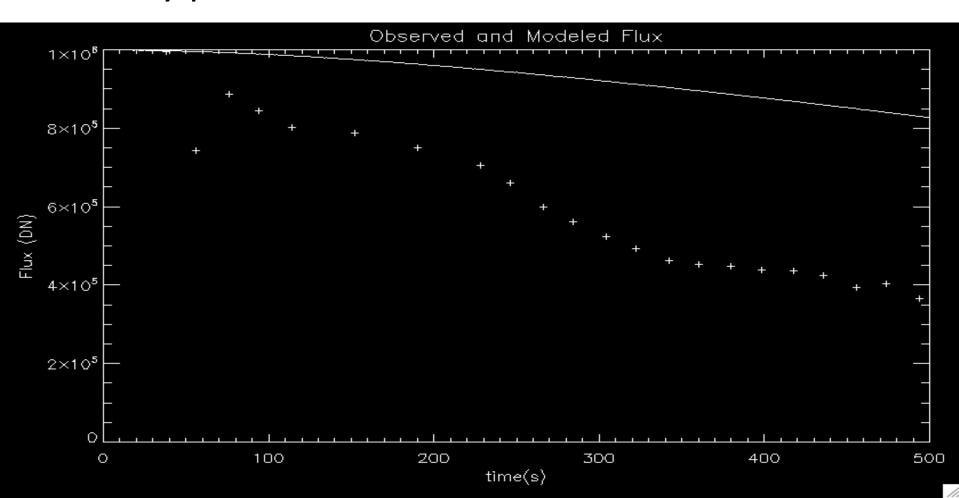
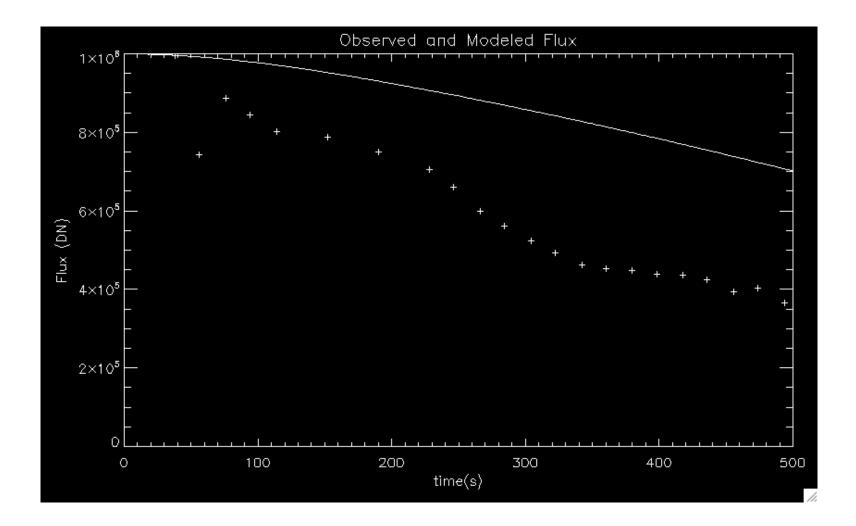
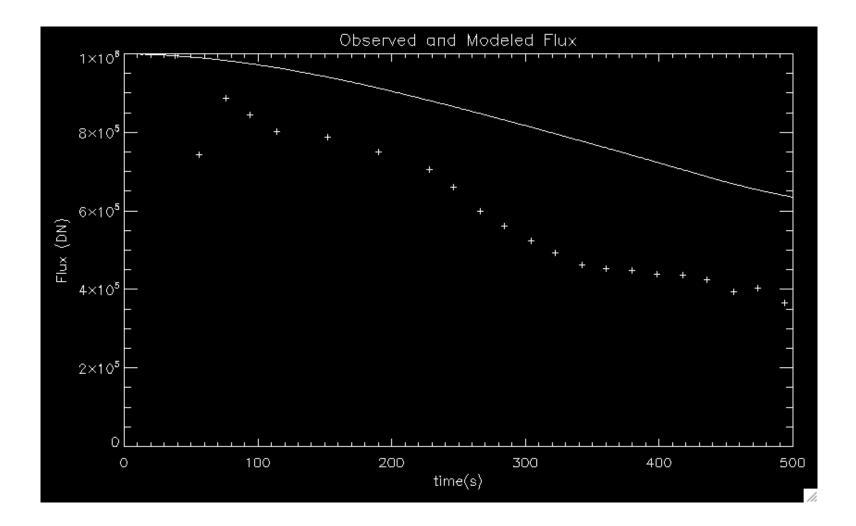


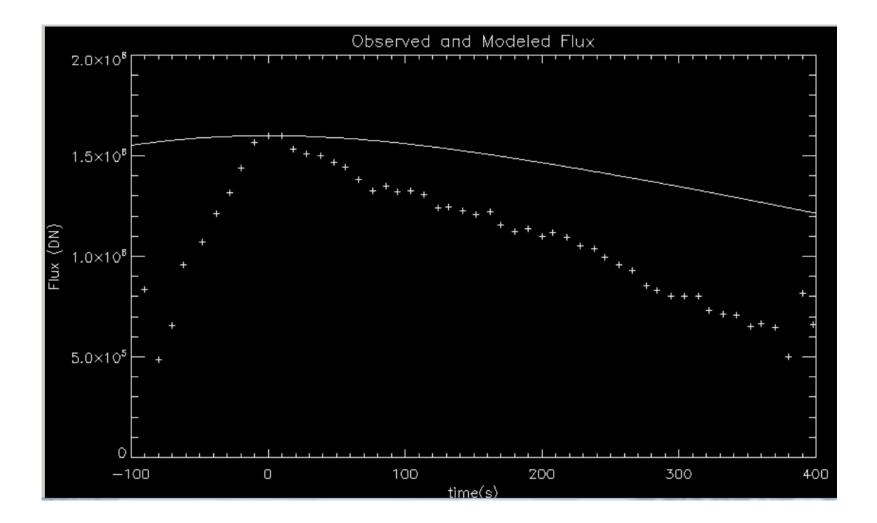
Fig. 2. Amoeba flowchart.

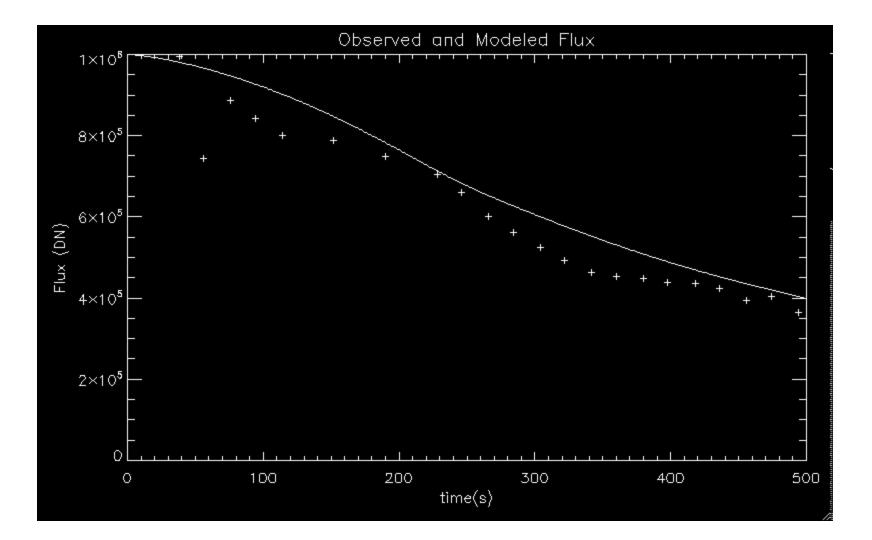
• Amoeba searches for the best fit line. In this early picture it is far off



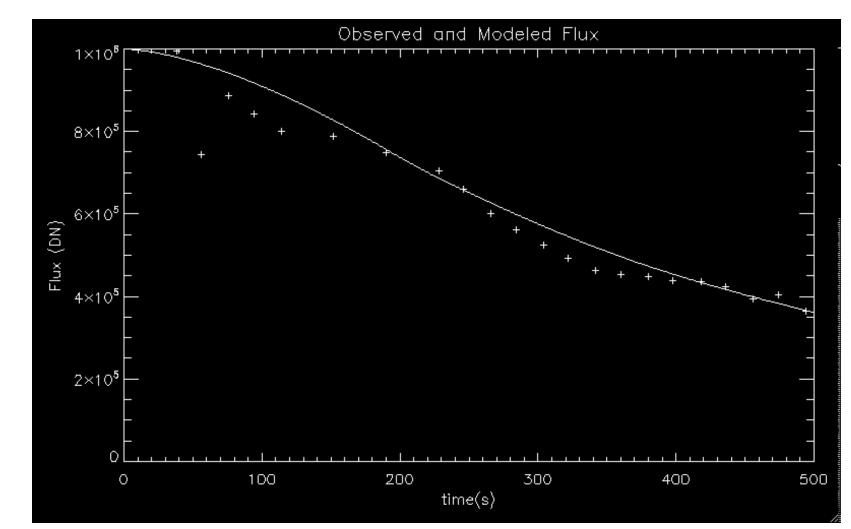




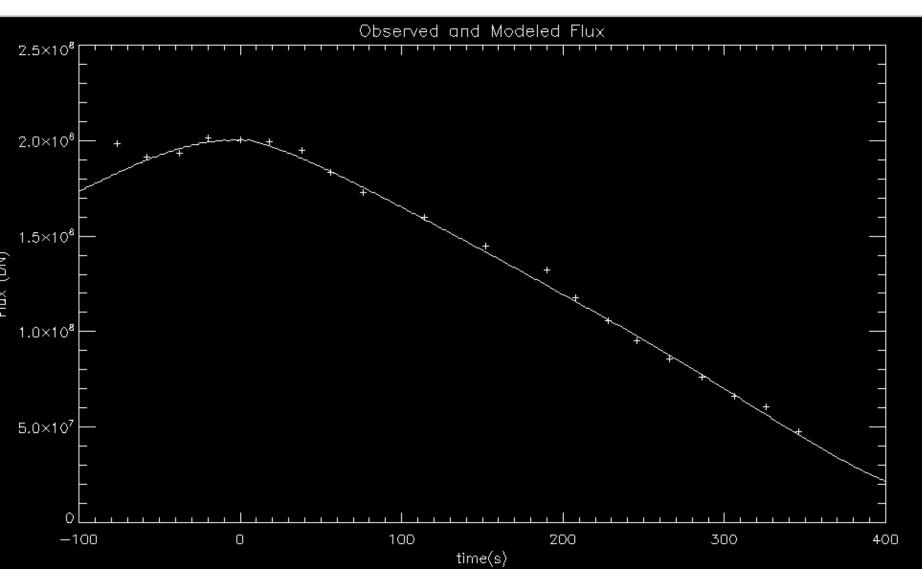




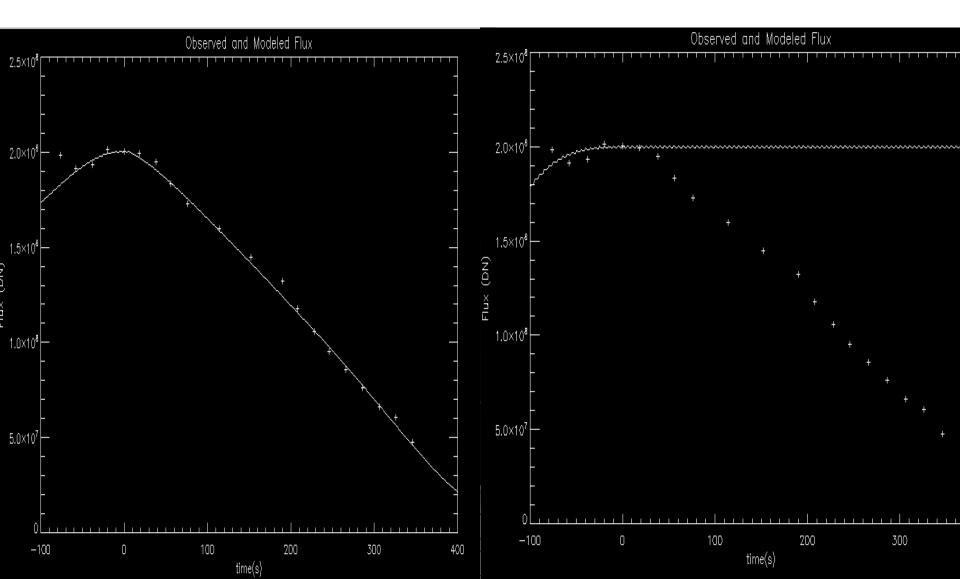
• Eventually it finds a pretty good fit.



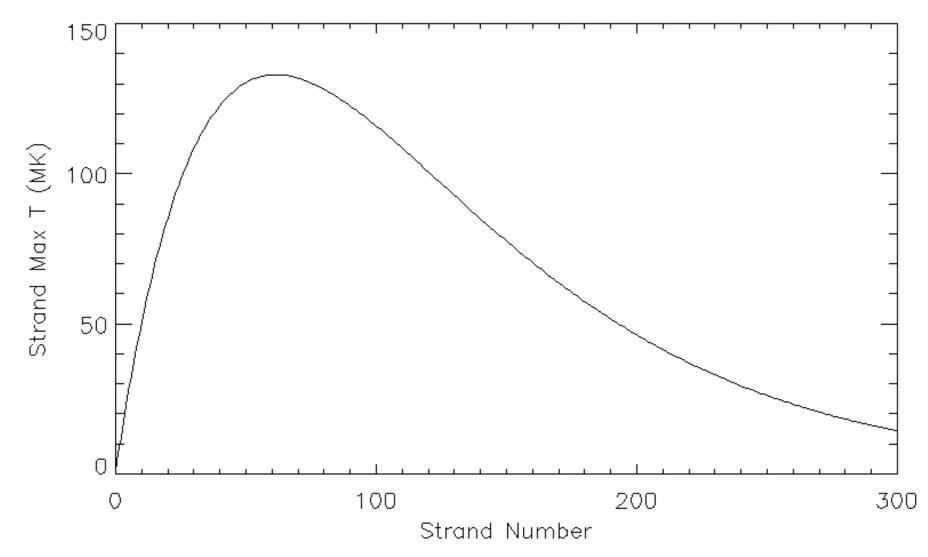
Best fit results to flare of September 23, 1998



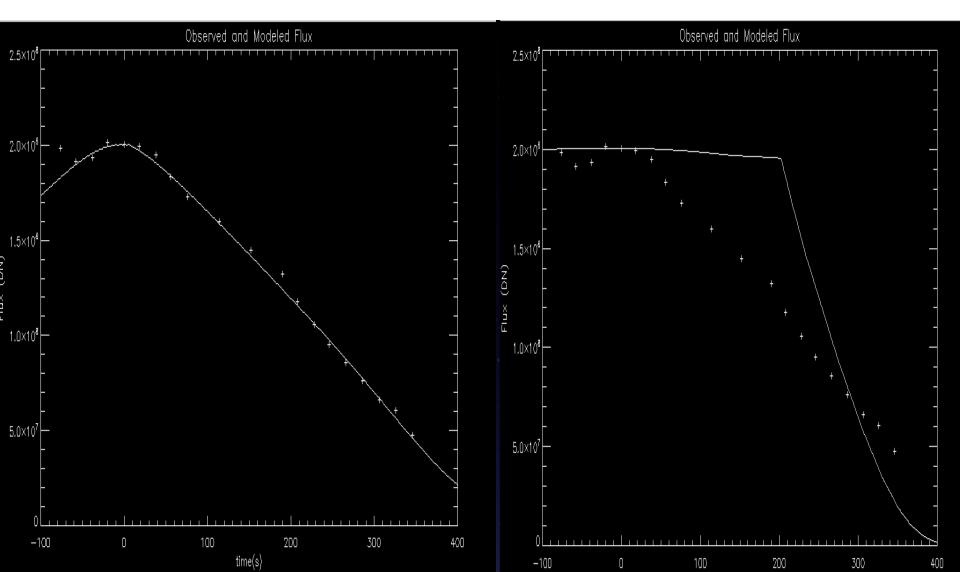
Comparison New to Old



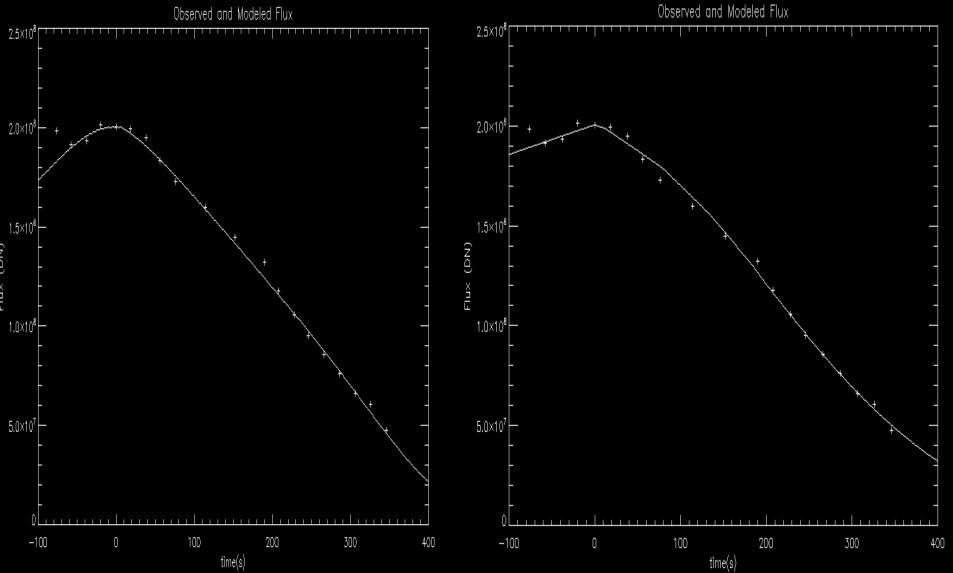
Amoeba Loop number vs peak temperature



Comparison Ours to Single Strand Model

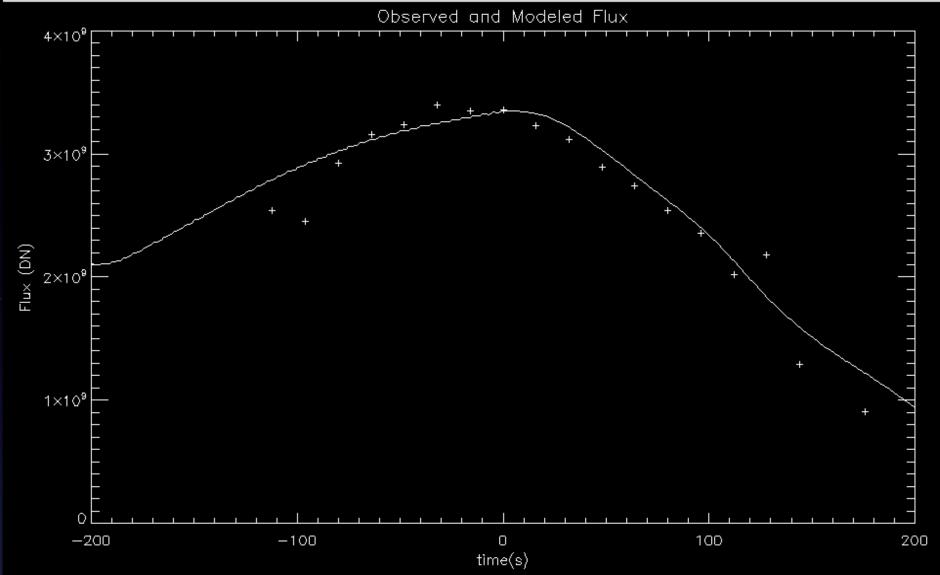


Single Strand Can Be Close but the Parameters Become Unreasonable

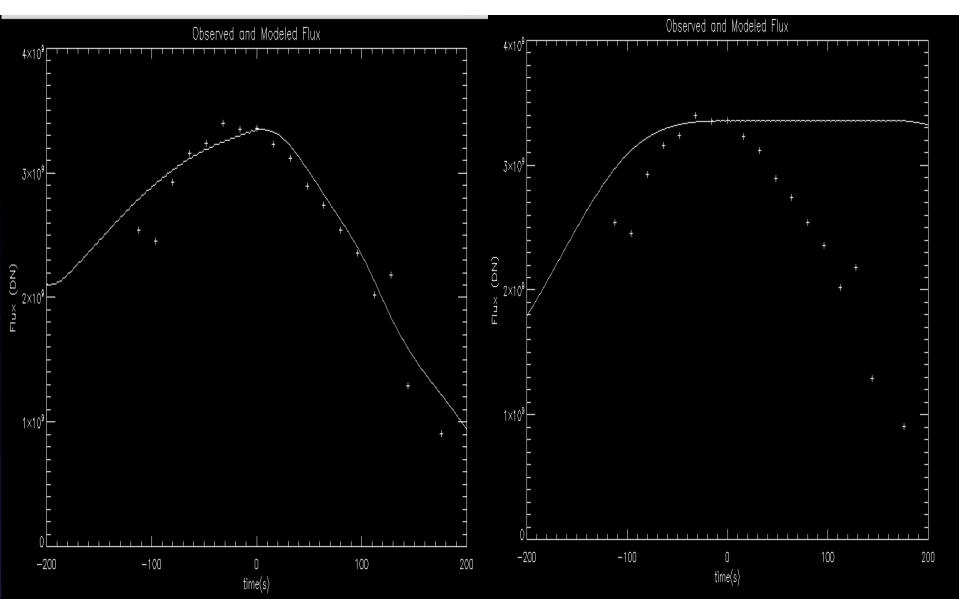


Best fit results to flare of December

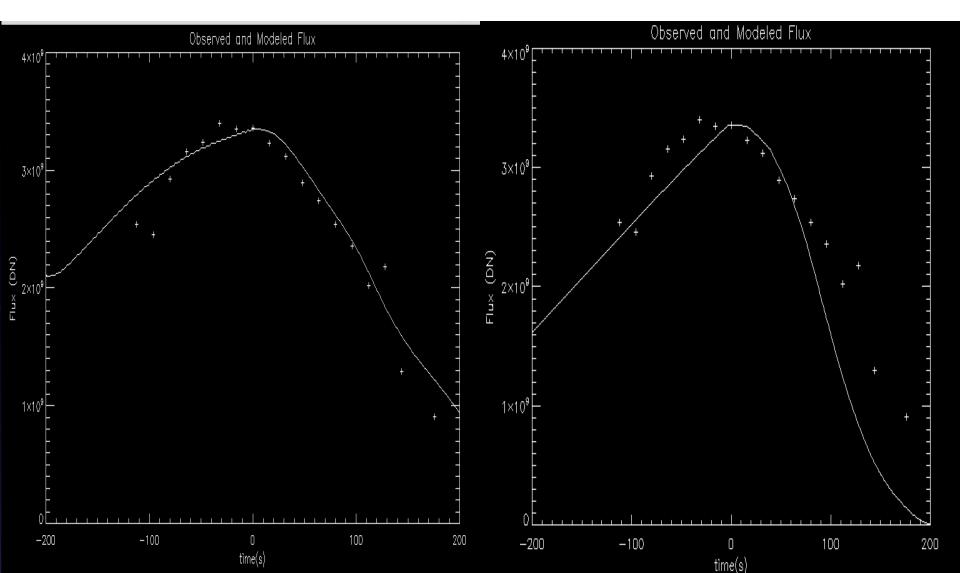
27th, 1999



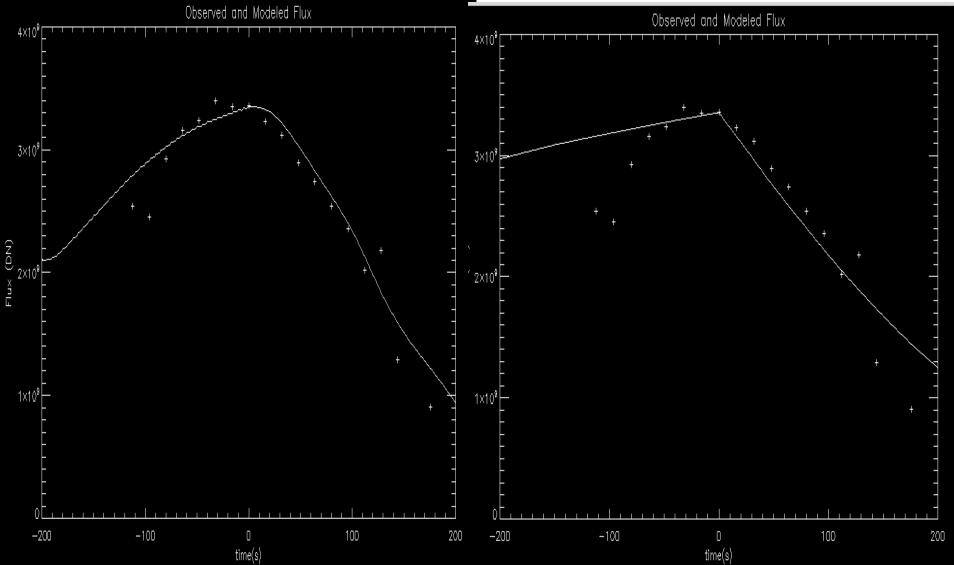
Comparison New to Old



Comparison Ours to Single Strand Model



Single Strand Can Be Close but the Parameters Become Unreasonable



The Future

 We have successfully improved our model so that it better fits the cooling of solar flare loops. In the future, we hope to improve our bounds on parameter space so that our fits can be even better and our model can be physically accurate.

